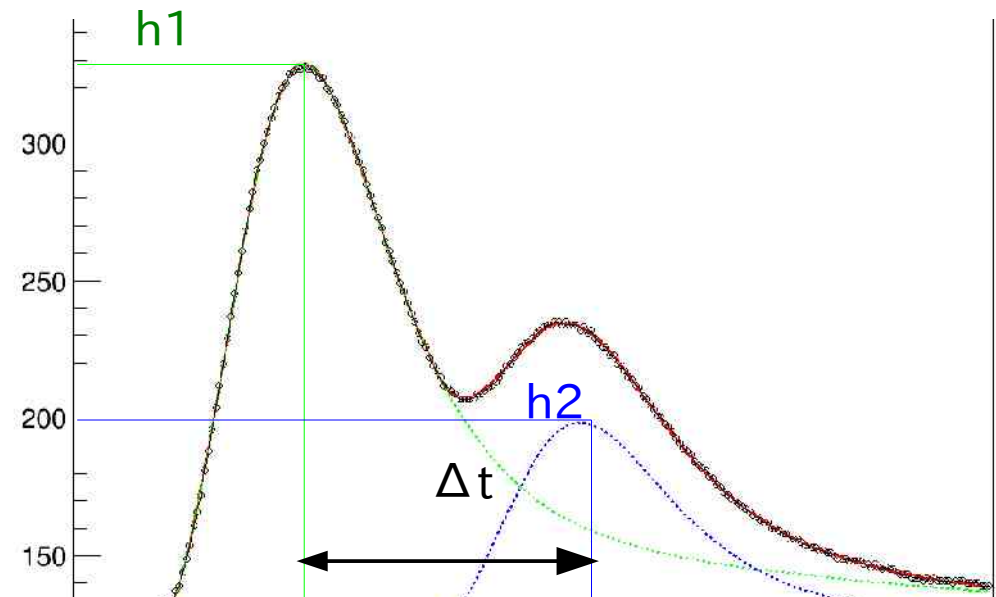
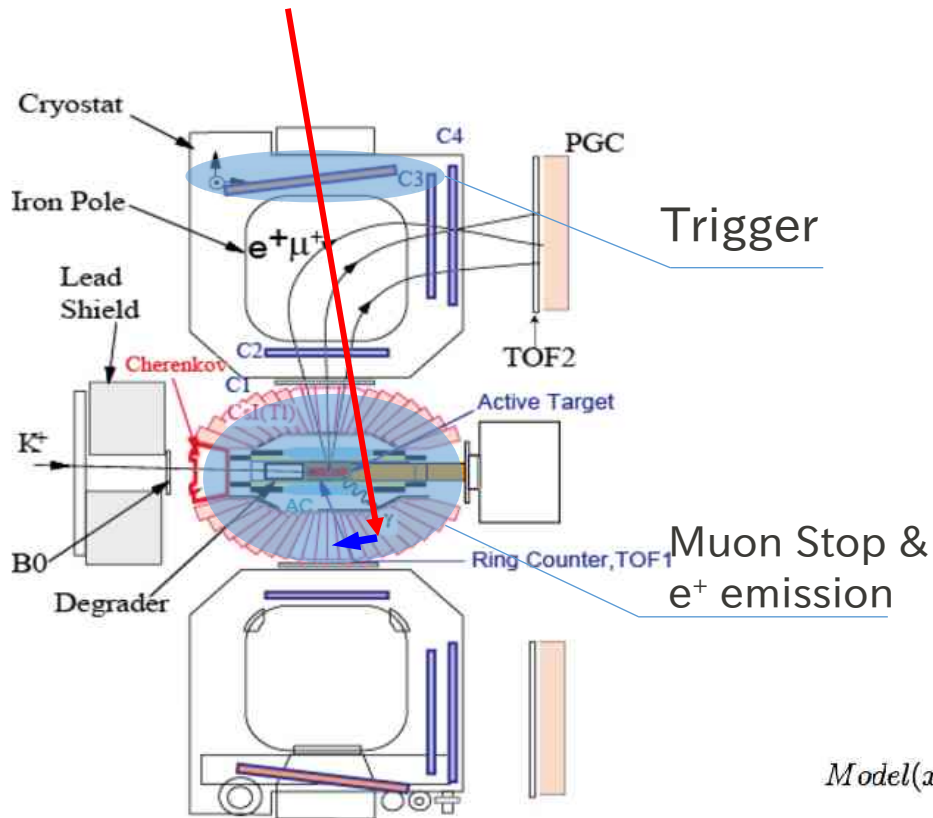


Calibration of the CsI(Tl) detector using Cosmic muons

2016.03.25

H. Ito



$$Model(x) = p_0 Freq \left[\frac{x - p_1 - p_2}{p_3} \right] \frac{1}{f(x_0)} \times$$

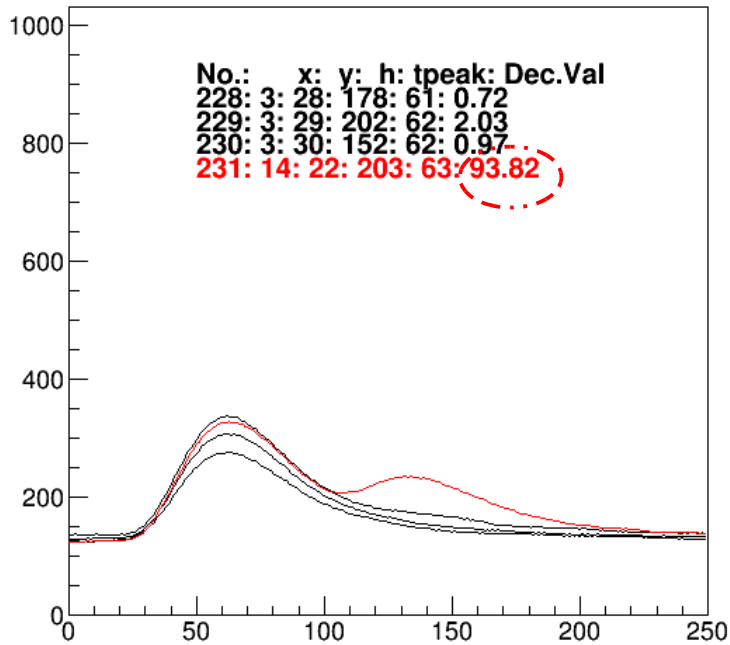
$$\left(\frac{x - p_1}{p_4} \exp \left[1 - \frac{x - p_1}{p_4} \right] + p_5 \frac{x - p_1}{p_4 + p_6} \exp \left[1 - \frac{x - p_1}{p_4 + p_6} \right] \right) \theta(x - p_1) + p_7$$

$$f(x_0) = \frac{\varepsilon \tau_1 (\varepsilon \tau_1 + \tau_2)}{\varepsilon \tau_1^2 + \tau_2^2} \exp \left(1 - \frac{\tau_1 (\varepsilon \tau_1 + \tau_2)}{\varepsilon \tau_1^2 + \tau_1^2} \right) + \frac{\tau_2 (\varepsilon \tau_1 + \tau_2)}{\varepsilon \tau_1^2 + \tau_2^2} \exp \left(1 - \frac{\tau_2 (\varepsilon \tau_1 + \tau_2)}{\varepsilon \tau_1^2 + \tau_2^2} \right)$$

$$Freq(x) = \frac{1}{\sqrt{2\pi}} \int_{-\infty}^x \exp(-t^2/2) dt$$

Analysis Step.

wf_228 (n=1791, run=4563 x=3 y=28)



Step. 1: Single waveform fitting

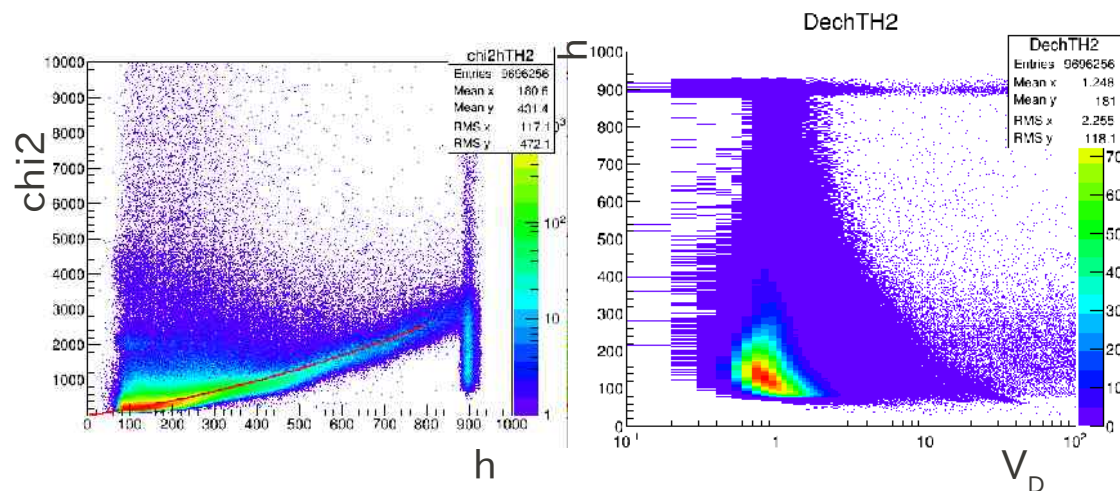
Step. 2: $V_D > 2$

Step. 3: Double waveform fitting

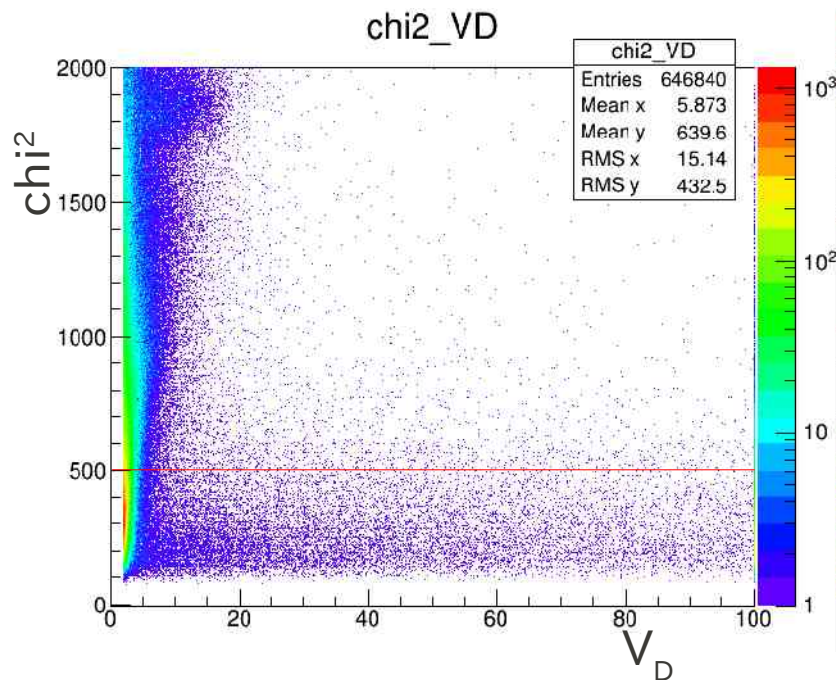
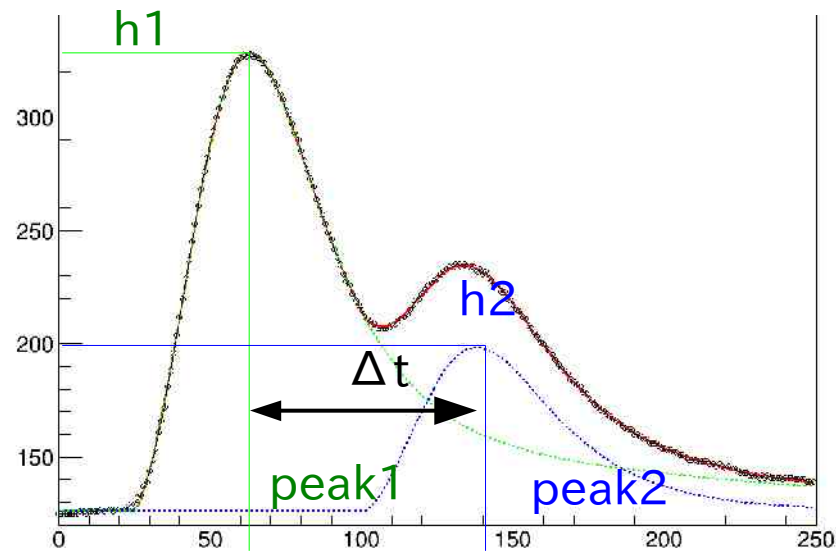
Step. 4: $\text{chi}^2 < 500$

Step. 5: 2D-hist (h2 vs. Δt)

Step. 6: Rejected Noise event



Analysis Step.



Step. 1: Single waveform fitting

Step. 2: $V_D > 2$

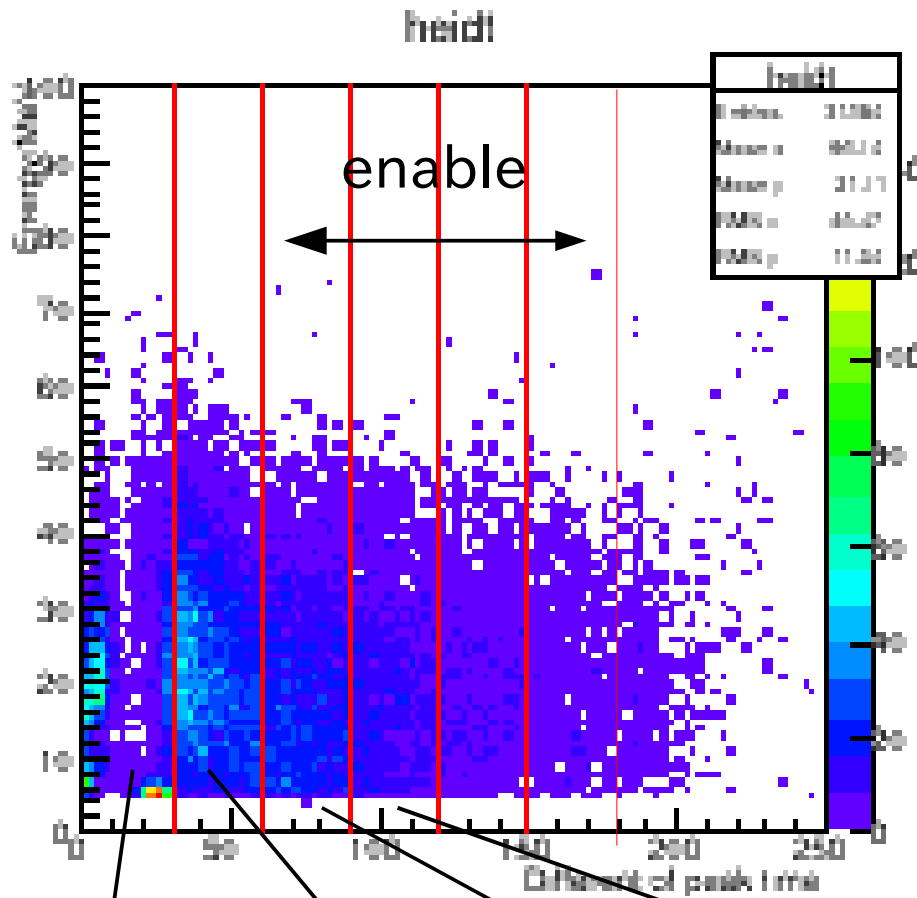
Step. 3: Double waveform fitting

Step. 4: $\text{chi}^2 < 500$

Step. 5: 2D-hist (h_2 vs. Δt)

Step. 6: Rejected Noise event

Analysis Step.



Step. 1: Single waveform fitting

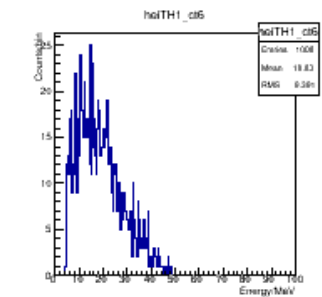
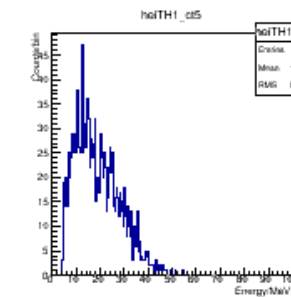
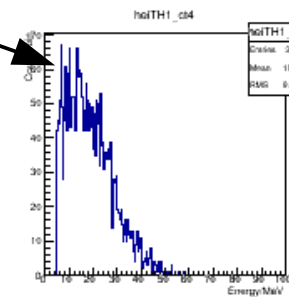
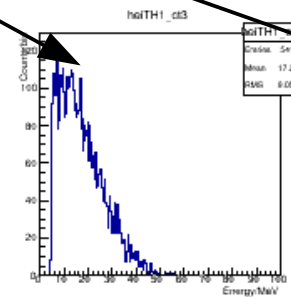
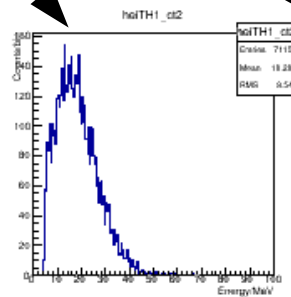
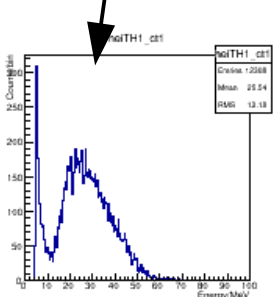
Step. 2: $V_D > 2$

Step. 3: Double waveform fitting

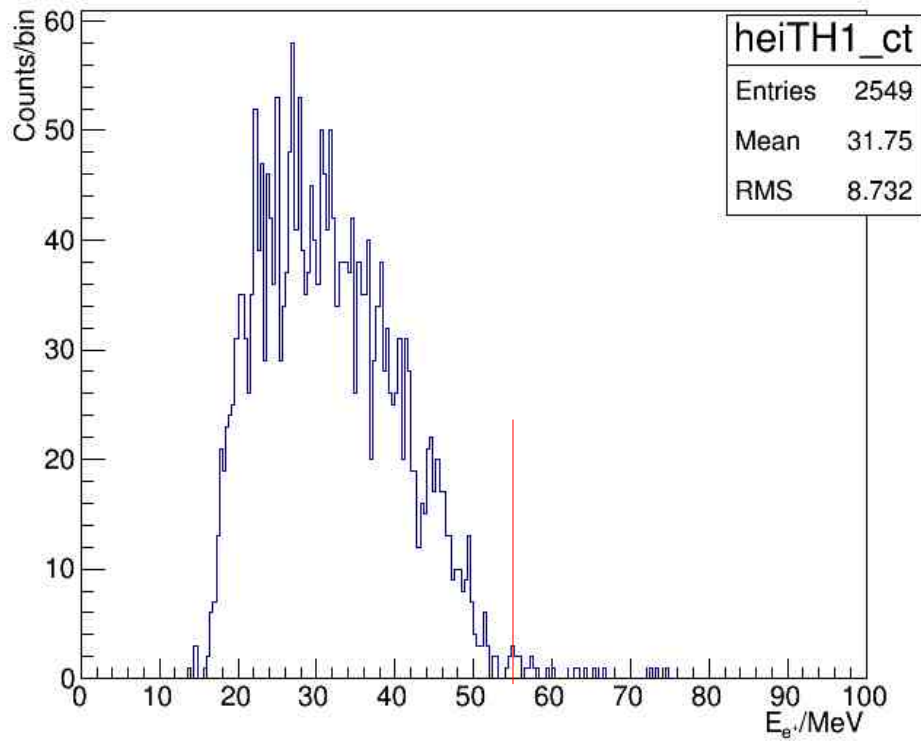
Step. 4: $\chi^2 < 500$

Step. 5: 2D-hist (h2 vs. Δt)

Step. 6: Rejected Noise event

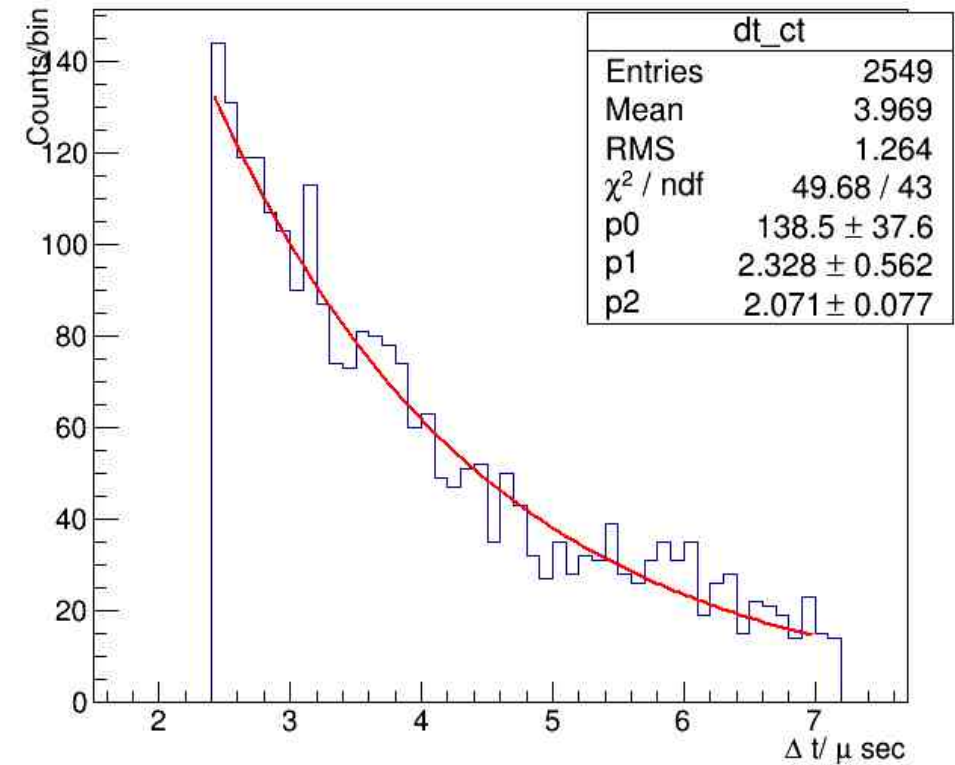


Result



E_{e^+} is an energy of e^+ decayed from μ , maximum 53 MeV.

It consists in comparison with calibration with Kmu2 (153 MeV).



Fitting $f(x) = p_0 \cdot \exp(-(x-p_1)/p_2)$
 p_2 = mean life time of μ is 2.07 ± 0.08 us.
 ($\tau_\mu = 2.19$ us in vacuum)

Next work

wf_726 (n=5263, run=4563 x=8 y=23)

